

Study of X-ray emission from the old open cluster, M67

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ABSTRACT

We present an X-ray analysis of a 4 Gyr old open cluster, M67, using archival *XMM–Newton* data. The aim of this study was to find new X-ray members of M67, and to use the updated member list for studying X-ray variability, and derive the X-ray luminosity functions (XLFs) of different stellar types and compare them with other star clusters of similar age. We report the detection of X-ray emission from 25 members of M67, with membership based primarily on their proper motion, of which one X-ray source is a new member. Supplementing this study with previous *ROSAT* and *Chandra* studies of M67, and using the most recent proper motion study by Vereshchagin et al., we have compiled a revised list of X-ray emitting members of M67 consisting of 43 stars. 16 of these are known RS CVn type binaries having orbital periods <10 d, and near-circular orbits, five are contact binaries with orbital periods <6 h, five are yellow and blue stragglers, two are Algol-type binaries, and one source is a cataclysmic variable. 14 members do not have any orbital information and cannot be classified. 14 of the X-ray sources detected do not have any optical counterpart down to a magnitude of $V \simeq 22$, and their membership is uncertain. Finally, we report the XLFs of RS CVn type and other types of stars in M67 and compare them with other open clusters of intermediate-to-old age.

Key words: stars: activity – binaries: general – blue stragglers – open clusters and associations: individual: M67 – X-rays: binaries.

1 INTRODUCTION

Open clusters are useful for studying coeval and comoving populations of stars within the Galactic disc. X-ray studies of stars in open clusters offer an insight into their coronal activity and/or accretion phenomena. As clusters age, the spin-down of stars causes X-ray emission to diminish in general, and thus, in the X-rays, revealing active coronae primarily from stars spun-up in binary systems (the age–rotation–activity correlation; e.g. Pallavicini 1989; Randich 1997; Gudel et al. 2004) or from systems undergoing accretion. Accordingly, X-rays from old (a few Gyr or older) open clusters are unique probes of magnetically active (RS CVn, BY Dra, W UMa, FK Com, Algol) and mass-transfer (CVs, L/HMXBs) binary systems within the Galactic disc (Belloni, Verbunt & Mathieu 1998; Verbunt 1999; van den Berg 2013).

Past studies of old open clusters, NGC 6791 (~8 Gyr; van den Berg et al. 2013), NGC 188 (~6 Gyr; Belloni et al. 1998), and M67 (Belloni, Verbunt & Schmitt 1993; Belloni, Verbunt & Mathieu 1998; van den Berg et al. 2004), reveal that RS CVns, CVs, and sub-giant stars dominate the X-ray emission, while peculiar objects

such as blue and yellow stragglers are rare and thus do not contribute much.

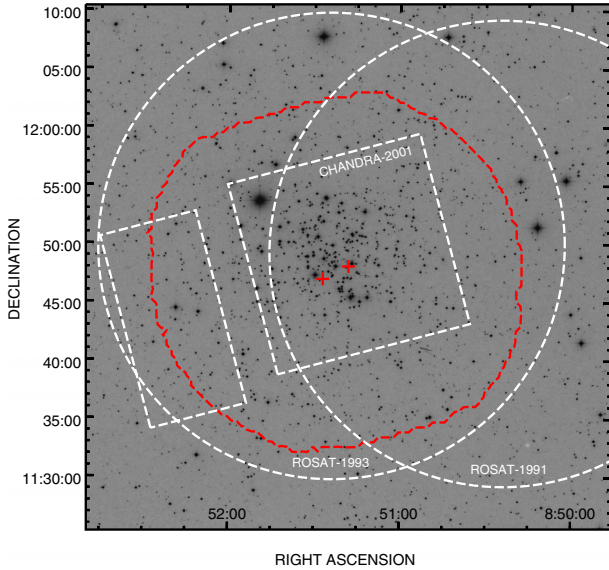
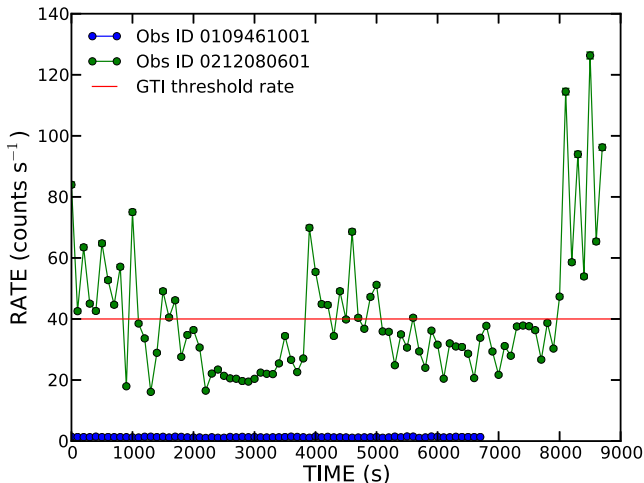
M67 is a 4.2 ± 0.6 Gyr-old open cluster at 850 ± 30 pc having a small reddening value ($E(B - V) \approx 0.04$; Sarajedini et al. 1999; Yadav et al. 2008). On account of the extensive optical data available for M67, this cluster is well suited for study in X-rays. Proper-motion studies of M67 to establish cluster membership have been carried out by several groups (Sanders 1977; Girard et al. 1989; Zhao et al. 1993; Yadav et al. 2008; Vereshchagin et al. 2014). In this paper, we use cluster membership information from Vereshchagin et al. (2014), a revised version of the Yadav et al. (2008) catalogue reaching down to $V \sim 22$ mag, containing 659 members. Two published X-rays studies of M67 exist. Belloni et al. (1993, 1998) presented *ROSAT* PSPC observations covering ~ 0.5 deg² of the M67 field, while van den Berg et al. (2004) analysed *Chandra* ACIS observations covering ~ 0.1 deg² but with a limiting flux of about 40 times lower than the *ROSAT* observations.

Here, we present the results of two *XMM–Newton* observations of M67 having fields of view and limiting flux intermediate to that afforded by the *ROSAT* and *Chandra* observations. The paper is organized as follows. In Section 2, we discuss the X-ray data and processing. The optical, *ROSAT* and *Chandra* counterparts of our X-ray sources are given in Section 3. The spectral hardness and variability analysis is in Section 4. Section 5 gives notes on individual

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Table 1. *XMM-Newton* data on M67.

Obs ID	0109461001	0212080601
Coordinates	08:51:26.99 11:46: 58.0	08:51:18.00 11:48:02.6
Start time (UT)	2001 Nov 20 23:56:16	08 May 2005 18:23:17
Stop time (UT)	2001 Nov 21 02:43:07	08 May 2005 22:23:28
Usable time (MOS1,MOS2,PN; ks)	9.42,9.42,6.80	5.70,5.70,5.70
Fiter, EPIC mode	Thin, Full frame	Medium, Full frame

**Figure 1.** A 45 arcmin \times 45 arcmin image of M67 from the Digital Sky Survey. The footprints of the *Chandra* ACIS observation (dashed white boxes; van den Berg et al. 2004), *ROSAT* observations (dashed white circles; Belloni et al. 1993, 1998), and the *XMM-Newton* observations used in this work (red; aimpoints marked by crosses) are overlaid.**Figure 2.** The > 10 keV PN light curves for the two *XMM-Newton* observations of M67, ID 0109461001 (blue) and ID 0212080601 (green). The latter has enormous soft proton flaring, and the count rate threshold of 40 counts per second used to define the GTI is shown with the red line.

classes of X-ray sources, and we conclude with a discussion in Section 6.

2 ARCHIVAL DATA AND DATA PROCESSING

2.1 X-ray data

We used archival data from two observations of M67 from the *XMM-Newton* Science Archive (XSA). Observation with ID 0109461001 was carried out in 2001 and ID 0212080601 in 2005. For both the observations, we used the European Photon Imaging Camera (EPIC, consisting of two MOS and one PN CCD arrays; Jansen et al. 2001; Strüder et al. 2001; Turner et al. 2001). The first observation used the thin filter while the second observation used the thick filter. Data were acquired in the full frame mode in both the cases. The observation details are given in Table 1. The merged *XMM-Newton* footprint for the two M67 observations is shown in Fig. 1. The footprints of the previously published *ROSAT* and *Chandra* observations of M67 are also shown. Although the 3*XMM* pipeline products (Rosen et al. 2015) are available, we manually reduced and inspected the data to use the latest calibration files and a better control over the filtering, reduction and source selection process. We downloaded observation data files (ODFs) from the XSA for further processing.

2.2 Data processing

The raw data were processed with the *SAS* 13.5.0 and *HEASOFT* 15.1 packages using the procedure outlined in the *XMM* ABC and *XMM* *SAS* guidebooks. In brief, we (i) generated the calibration information file using *cifbuild* task and downloaded the relevant calibration files from the *XMM* calibration archive, (ii) ran *odfingest* to compile housekeeping information on the ODFs, (iii) applied the calibration with the *emchain* and *epchain* tasks, (iv) filtered the MOS and PN event lists by selecting only the good events (*XMMEA_EM* and *XMMEA_EP*, respectively). We then inspected the PN light curve for photons with energies > 10 keV and found no evidence for soft-proton flaring in the ID 0109461001, but ID 0212080601 was contaminated by a significant flaring. To mitigate the flaring background in ID 0212080601, we chose a good time interval (GTI) where the PN count rate for energies > 10 keV was more than 40 counts s^{-1} . The > 10 keV PN light curves for ID 0109461001 and ID 0212080601 along with the GTI count rate threshold for the latter are shown in Fig. 2. Using the GTI removes a large fraction of the major flaring events and reduces the background in the image by a factor of 4. Note that the count rate threshold chosen for GTI is much larger than the standard 0.4 counts s^{-1} , but is acceptable for the detection of point sources.

2.3 X-ray images

We produced images in three energy bands, a soft (S_X) band (0.2–0.5 keV), a medium (M_X) band (0.5–2.0 keV) and a hard (H_X)

band (2.0–7.0 keV) band, for all three EPIC detectors using the `evselect` and `emosaic` tasks with the filtered MOS and PN event lists from Section 2.2 as input. Total energy band (0.2–7.0 keV) images for each detector were also produced. Although *XMM-Newton* is sensitive up to 10 keV, we restricted our images to <7 keV in order to enable easy comparison with the *Chandra* observation of M67 (van den Berg et al. 2004). During the imaging process, only events with (i) the *FLAG* keyword set to zero, (ii) *PATTERN* keyword less than or equal to 4 for PN and (iii) *PATTERN* keyword equal to zero for PN in the soft band were selected. The motivation for the latter is mainly the rejection of noise at the extremities of the PN detector along the detector *Y*-direction. This stricter requirement for *PATTERN* = 0 for PN in the soft band is also implemented in the *2XMM* catalogue (Watson et al. 2009), for example. Image binning was so chosen as to produce 600×600 pixel 4.35 arcsec pix^{-1} images.

2.4 Source detection

We used the `edetect_chain` script to simultaneously search for sources in the nine images (3 bands × 3 detectors) for each observation. The subroutine `edetect_chain` calls a series of tasks sequentially: (i) `expmap` to calculate the exposure map for the input images using the attitude, vignetting, exposure, bad pixel information, (ii) `emask` produced a detector mask based on the exposure, (iii) `eboxdetect` was used in local mode to search for sources above a maximum likelihood (ML) of 5 simultaneously in all the input images in 5×5 pix boxes (pixels surrounding the boxes are used for background estimation), (iv) `esplinemap` removed the sources detected by `eboxdetect` from the input images to produce smoothed background maps through spline fits to the residual images, (v) `eboxdetect` was used in the map mode to find sources similar to the earlier `eboxdetect` call, except this time the background map was used, (vi) `emldetect` used the source locations from map mode `eboxdetect` to perform simultaneous ML point spread function fits (Craddace, Hasinger & Schmitt 1988) to all input images and determine source parameters such as total counts, count rate, hardness ratios, etc. For `emldetect` we specified an ML detection threshold of six (similar to the *2XMM* catalogue; this corresponds to a 0.2 per cent probability for Poissonian noise fluctuation to have caused the observed source counts). The hardness ratios (discussed in detail in Section 4.1) were calculated for the detected sources based on their count rate in different bands, and these are especially useful for characterizing weak sources for which spectral analysis was not possible.

We combined the single-band and single-detector images into a single mosaic image using the `emosaic` task, for each observation. We inspected the mosaics by eye and further used the hardness ratios of detected sources as well as the knowledge of their optical counterparts to reject one false positive in ID 0109461001 from the `emldetect` source list. 72 sources are thus detected in ID 0109461001 and 32 in ID 0212080601. 29 sources are common between the two observations and three are unique detections in ID 0212080601. The conversion of count rates to fluxes is described in Section 4. The minimum 0.2–7 keV flux of detected sources in the two observations is 1.6×10^{-15} and $4.8 \times 10^{-15} \text{ erg cm}^{-2} \text{ s}^{-1}$, respectively. These correspond to X-ray luminosities of 1.4×10^{29} and $4.2 \times 10^{29} \text{ erg s}^{-1}$, respectively, at the assumed distance of 850 pc to M67. The images in the total energy band were subjected to `esensmap` to find the median count rate corresponding to likelihood threshold of six. Converting these to fluxes we get the median flux detection thresholds of 8.9×10^{-15} and $3.0 \times 10^{-14} \text{ erg}$

$\text{cm}^{-2} \text{ s}^{-1}$ in the 0.2–7.0 keV energy range for ID 0109461001 and ID 0212080601, respectively. Our final X-ray source catalogue is given in Table 2, which gives source parameters for the PN detector with count rates and total counts in the 0.2–7.0 keV band. The 1σ source position uncertainties range between 0.5 arcsec and 4.0 arcsec, the median uncertainty being 1.5 arcsec.

3 IDENTIFICATION OF X-RAY SOURCES

In order to identify the X-ray sources found in both *XMM-Newton* observations of M67 with known classes of astronomical sources, we searched for optical and previously known X-ray counterparts. Information from SIMBAD and infrared data from *WISE* were used as necessary in order to accurately classify sources (especially in case of AGN or foreground stars). All identifications were checked by eye. Counterparts from Yadav et al. (2008) were used to establish membership probabilities. Below, we describe the identification process, and give simple estimates of chance identification and expected background X-ray sources.

3.1 X-ray counterparts from *ROSAT* and *Chandra*

Previous observations with *ROSAT* (Belloni et al. 1993, 1998) and *Chandra* (van den Berg et al. 2004) were with flux detection thresholds of 9×10^{-15} and $2 \times 10^{-16} \text{ erg cm}^{-2} \text{ s}^{-1}$, respectively (see Fig. 1 for coverage). A comparison of our X-ray sources with those in the published catalogues, shows that, out of 61 X-ray sources detected by Belloni et al. (1998) with the *ROSAT* PSPC, 44 are in the *XMM* field of view (fov), of which, 40 have one *XMM* counterpart within the positional error. van den Berg et al. (2004) detected 158 X-ray sources using the *Chandra* ACIS-I and ACIS-S detectors, of which 153 *Chandra* sources lie within the *XMM* fov, and 41 of these have a probable *XMM* counterpart. The X-ray counterparts are listed in columns (8) and (9) of Table 2.

3.2 Optical counterparts

The most comprehensive and sensitive survey of M67 is the ESO Imaging Survey Pre-FLAMES (EIS; Momany et al. 2001), which contains objects with $11 \leq V \leq 23$. We cross-matched our X-ray source list with the EIS catalogue to find optical counterparts. The error in the optical source positions was taken to be 1 arcsec, and this was added in quadrature to the X-ray source position uncertainties. This procedure gave unique EIS counterparts for 50 *XMM* sources. For sources absent in EIS, we searched for matches in the SDSS and Yadav et al. (2008) catalogues, thus finding six additional sources with optical counterparts. In this paper, we refer to the EIS and Yadav et al. sources with the prefixes ‘E’ and ‘Y’ respectively, e.g. the optical counterparts of NX1 and NX5 are E1740 and Y1289, respectively. The optical counterparts, their *V* magnitudes, *B* – *V* colours and the distance between the X-ray source and the optical counterparts are listed in columns (14), (10), (11), and (12) of Table 2, respectively. Optical magnitudes and/or colours unavailable in the EIS or Yadav et al. catalogues were estimated from other catalogues in Vizier Catalog Service,¹ wherever applicable. For SDSS counterparts, we converted *u,g,r* magnitudes to *V* magnitudes and *B* – *V* colours using the transformations for stars from Jester

¹ <http://vizier.u-strasbg.fr/>

Table 2. Summary of EPIC detections in the field of M67. For each source we give the *XMM-Newton* source ID (NX), coordinates (α_{J2000} , δ_{J2000}), uncertainty in the coordinates ($\delta_{\alpha, \beta}$), PN total counts (counts) and count rate (ctr) in the total energy band (0.2–7 keV), hardness ratios (HR), *ROSAT* and *Chandra* counterparts (RX and CX from Belloni et al. and van den Berg et al., respectively), the distance between the EPIC source position and the optical counterpart (d_{XO}), membership in M67 (Y/N: Yes/No, from Yadav et al.; n: no from multiwavelength data) and comments. ‘N!’ denotes sources that are not proper-motion members according to Yadav et al. and Vereshchagin et al., but show other evidences of membership (e.g. spectral fitting, literature), thus leading to their final classification as members of M67. The superscript ^a on the source ID denotes that the corresponding counts, count rate, and hardness ratios are from the MOS1 detector due to the source being in one of the gaps of the PN CCD.

(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)	(9)	(10)	(11)	(12)	(13)	(14)	
NX	α_{J2000} (^h , ^m , ^s)	δ_{J2000} ([°] , ['] , ^{''})	$\delta_{\alpha, \beta}$ (arcsec)	counts	ctr (c ks ^{−1})	HR1	HR2	RX	CX	V	B − V	d_{XO} (arcsec)	Mem	Identification
Sources detected in Obs ID 0109461001														
1	08:51:23.22	+11:48:25.9	0.4	291(19)	51(3)	0.63 ± 0.05	−0.84 ± 0.04	11	5	14.32	0.71	1.6	Y	E1740
2	08:51:38.40	+11:49:05.4	0.5	286(18)	53(3)	0.59 ± 0.06	−0.65 ± 0.05		2	20.79	0.15	1.2	n	E510; Gal, z~0.6
3	08:50:57.09	+11:46:07.6	0.5	208(19)	94(8)	0.39 ± 0.09	−0.50 ± 0.09	15	4	23.27	1.36	0.2	n	SDSSJ0850.95; AGN2, z~0.5
4	08:51:13.30	+11:51:39.5	0.6	169(14)	37(3)	0.68 ± 0.07	−0.91 ± 0.05	8	1	13.57	1.01	1.6	Y	E1708
5	08:51:23.78	+11:49:49.3	0.5	151(13)	28(2)	0.82 ± 0.06	−0.90 ± 0.05	10	6	11.44	0.90	0.1	Y	Y1289
6 ^a	08:51:18.64	+11:47:02.5	0.6	125(12)	14(1)	0.70 ± 0.08	−0.74 ± 0.07	13	9	12.67	0.74	1.1	Y	E1726
7	08:51:28.30	+11:41:27.4	0.7	123(12)	26(2)	0.55 ± 0.09	−0.56 ± 0.10		12	22.40	0.12	1.6	N	E2024
8	08:51:07.21	+11:53:02.0	0.6	118(16)	53(7)	0.70 ± 0.11	−0.65 ± 0.13	7	10	12.68	0.69	0.7	Y	E1687
9	08:51:51.35	+11:35:47.5	0.8	72(10)	28(3)	0.39 ± 0.13	−0.64 ± 0.14	59		20.99	0.66	2.0	n	E1907; QSO1
10	08:51:37.79	+11:50:56.6	0.8	68(9)	13(1)	0.58 ± 0.12	−0.90 ± 0.08	40	16	13.64	0.45	1.4	y	E1791; AH Cnc (F7V)
11	08:51:49.93	+11:49:52.0	0.9	67(9)	14(2)	0.39 ± 0.14	−0.68 ± 0.13	43	20	20.98	−0.05	0.9	N	E2757
12	08:50:52.70	+11:47:44.4	1.1	62(9)	19(2)	0.22 ± 0.16	−0.40 ± 0.17	49	21	21.86	0.27	1.4	n	E2183; AGN2
13	08:51:33.02	+11:34:11.9	1.3	62(9)	26(4)	0.74 ± 0.11	−0.77 ± 0.13	60		21.03	0.56	1.7	n	E2593; Gal
14	08:51:04.52	+11:55:28.1	0.8	61(9)	20(3)	0.72 ± 0.13	−0.45 ± 0.14	34	14	21.99	0.81	3.1	n	E2823; Gal
15	08:50:37.97	+11:47:09.3	1.2	54(9)	24(3)	0.72 ± 0.16	−0.18 ± 0.17	12		22.94	2.01	2.3	n	SDSSJ0850.63; AGN2
16	08:51:21.63	+11:52:37.5	0.9	52(8)	11(1)	0.64 ± 0.13	−1.00 ± 0.08	37	24	11.26	0.66	1.9	Y	Y1476
17	08:52:16.79	+11:48:31.0	1.1	50(8)	20(3)	0.69 ± 0.13	−0.83 ± 0.12	47	7	14.44	0.80	1.3	Y	E1857
18	08:51:03.86	+11:46:29.4	1.6	50(11)	20(4)	−0.04 ± 0.29	0.05 ± 0.30	14	13	21.21	0.36	1.0	N	E2153
19 ^a	08:51:37.27	+11:59:02.5	1.1	49(8)	12(2)	0.51 ± 0.14	−1.00 ± 0.08	31		11.62	0.06	1.3	N	E1787
20	08:51:28.12	+11:49:26.6	0.9	48(8)	8(1)	0.52 ± 0.15	−0.93 ± 0.10	45	19	12.84	0.41	1.4	Y	E1756
21	08:51:04.76	+11:45:56.3	1.1	47(8)	12(2)	0.87 ± 0.10	−1.00 ± 0.06		23	13.58	0.51	1.6	Y	E1569
22	08:51:17.71	+11:44:30.4	1.4	43(7)	8(1)	0.71 ± 0.13	−0.99 ± 0.05	17	15	15.49	0.84	0.8	Y	E1175
23	08:51:46.08	+11:49:49.9	0.9	42(7)	9(1)	0.62 ± 0.16	−0.67 ± 0.16		22	22.12	1.20	1.3	n	E2756; Gal
24	08:51:50.15	+11:46:05.9	1.4	41(7)	8(1)	0.61 ± 0.15	−1.00 ± 0.10	52	47	10.70	0.96	1.3	N!	Y892
25	08:52:13.52	+11:46:21.1	1.7	39(8)	13(2)	0.72 ± 0.16	−0.81 ± 0.16	51		21.95	0.50	1.2	N	E2147
26	08:51:46.56	+11:52:01.7	1.2	36(7)	8(1)	0.38 ± 0.20	−0.60 ± 0.18		33	20.74	0.18	1.4	N	E600
27	08:51:07.92	+11:49:54.4	1.2	33(7)	7(1)	0.87 ± 0.13	−0.88 ± 0.13	42	17	15.68	0.86	2.5	Y	E2759
28	08:52:10.42	+11:47:12.7	1.6	33(7)	10(2)	0.41 ± 0.20	−0.55 ± 0.26			22.08	0.10	2.4	n	E2722; Gal
29	08:51:19.59	+11:52:10.5	1.2	32(10)	12(3)	0.66 ± 0.27	−0.89 ± 0.18	38	48	14.00	0.60	2.0	y	E1732; HX Cnc (G1V)
30	08:50:54.97	+11:56:50.3	1.5	28(6)	11(2)	0.47 ± 0.20	−1.00 ± 0.24		58	15.90	1.04	0.7	Y	E763
31	08:52:07.40	+11:52:07.1	1.4	27(6)	8(2)	−0.04 ± 0.24	−0.91 ± 0.23			21.30	0.77	0.6	n	E2952; Fore* (~M7V)
32	08:51:35.76	+11:47:59.7	1.1	25(6)	4(1)	0.77 ± 0.31	0.16 ± 0.26	48	31	22.38	0.34	1.7	n	E2190; Gal
33	08:51:36.03	+11:46:31.9	1.4	25(6)	4(1)	0.45 ± 0.24	−0.93 ± 0.15	50	49	12.79	0.63	2.0	Y	E1784
34	08:50:58.74	+11:51:36.3	1.4	25(6)	7(1)	0.29 ± 0.27	−0.61 ± 0.28	39	32	20.07	0.23	3.1	N	E1268
35	08:51:30.61	+11:50:45.6	1.7	24(6)	4(1)	0.56 ± 0.23	−0.84 ± 0.20		29					
36	08:51:31.43	+11:45:50.8	1.6	24(11)	10(5)	1.00 ± 0.71	0.14 ± 0.46	53	36	12.74	0.50	2.4	Y	E1581
37	08:51:20.83	+11:53:27.2	1.7	24(9)	13(5)	0.83 ± 0.31	−0.62 ± 0.34	4	3			0.8	N!	E3205
38	08:51:52.93	+11:36:54.7	1.5	24(6)	8(2)	0.59 ± 0.22	−0.55 ± 0.26			21.91	0.85	2.0	N	E1929
39	08:52:07.44	+11:41:13.2	1.2	24(6)	8(2)	−0.29 ± 0.28	−0.07 ± 0.40	55		21.52	0.15	1.8	N	E2018
40	08:51:11.60	+11:50:29.5	1.5	23(6)	5(1)	1.00 ± 0.15	−0.37 ± 0.24							
41	08:51:30.47	+11:33:39.8	1.8	23(6)	10(2)	0.41 ± 0.23	−0.89 ± 0.25			22.07	0.00	2.4	n	E1866; AGN1
42	08:51:27.27	+11:46:55.1	1.5	23(6)	3(1)	−0.50 ± 0.33	0.23 ± 0.46	16	57	21.75	0.18	2.3	N!	E2164, EU Cnc
43	08:51:04.66	+11:42:19.3	2.2	22(6)	6(1)	0.53 ± 0.24	−0.91 ± 0.18							
44	08:51:27.13	+11:45:02.4	1.3	22(7)	5(1)	0.64 ± 0.26	−0.87 ± 0.29			22.83		3.9	n	E3674; Gal
45	08:51:24.69	+11:48:59.8	2.4	22(7)	3(1)	0.56 ± 0.29	−0.67 ± 0.27	46		13.24	0.60	3.9	Y	Y1212
46	08:51:16.99	+11:48:22.8	2.2	20(6)	3(1)	0.72 ± 0.27	−0.99 ± 0.15							
47	08:50:32.58	+11:53:54.3	3.1	20(6)	10(3)	1.00 ± 0.21	−0.42 ± 0.32							
48	08:50:57.55	+11:52:55.2	1.8	19(5)	6(1)	0.83 ± 0.25	−0.05 ± 0.30	36						
49	08:51:09.00	+11:50:44.4	1.6	19(5)	4(1)	−0.44 ± 0.80	0.88 ± 0.20		39	20.62	1.40	1.6	n	SDSSJ0851.14; Gal, z~0.4
50	08:51:06.38	+11:37:53.5	1.4	19(6)	19(7)	0.74 ± 0.26	−0.37 ± 0.40			20.36	0.27	2.3	N	E111
51	08:51:57.78	+11:45:46.1	1.8	18(5)	4(1)	0.81 ± 0.23	−0.66 ± 0.25			15.59	0.71	1.6	N	E1189
52	08:50:53.19	+11:59:24.6	2.6	17(5)	9(3)	0.35 ± 0.41	0.15 ± 0.38			22.17	0.00	2.6	n	E2436; QSO1, z~1
53	08:51:26.71	+11:56:10.5	1.8	16(5)	5(1)	1.00 ± 0.27	−0.79 ± 0.24		61	15.90	1.07	2.0	N!	E742
54	08:51:27.02	+11:55:23.9	1.5	15(5)	4(1)	1.00 ± 0.29	0.44 ± 0.31	32	8	23.46	2.36	1.2	n	SDSSJ0851.45; AGN2, z~0.8
55	08:51:44.01	+11:48:49.7	1.8	15(5)	3(1)	0.93 ± 0.17	−0.77 ± 0.29		70	22.79	0.03	0.9	n	E2208; Gal
56	08:51:36.30	+11:44:43.8	2.3	15(5)	3(1)	0.56 ± 0.33	−1.00 ± 0.32		40					
57	08:52:05.18	+11:50:42.9	2.4	14(5)	4(1)	1.00 ± 0.21	0.21 ± 0.36							
58	08:51:38.40	+11:54:13.8	2.2	14(5)	3(1)	0.32 ± 0.34	−0.46 ± 0.41		71,149	21.42	0.57	0.9	N	E2807
59	08:50:56.41	+11:37:31.2	1.4	14(5)	6(2)	0.78 ± 0.42	0.28 ± 0.34			22.31		2.3	n	E3586; QSO2, z~0.7
60	08:51:47.73	+11:39:54.6	3.3	14(5)	4(1)	−0.01 ± 0.36	−0.81 ± 0.38							
61	08:51:38.72	+11:58:21.3	2.3	13(4)	4(1)	0.08 ± 0.39	−0.51 ± 0.44							
62	08:51:04.43	+11:41:51.6	1.9	12(4)	3(1)	0.58 ± 0.30	−1.00 ± 0.19	54		19.36	1.66	2.5	Y	E1498

Table 2 – *continued*

(1)	(2)	(3)	(4)	(5)	(6)	(7)		(8)	(9)	(10)	(11)	(12)	(13)	(14)
NX	α_{J2000} (^h , ^m , ^s)	δ_{J2000} ([°] , ['] , ^{''})	$\delta_{\alpha,\beta}$ (arcsec)	counts	ctr (c ks ^{−1})	HR1	HR2	RX	CX	V (mag)	B − V (mag)	d_{XO} (arcsec)	Mem	Identification
Sources detected in Obs ID 0109461001														
63	08:51:01.48	+11:49:26.9	2.1	12(4)	3(1)	0.28 ± 0.36	−0.83 ± 0.34	44	68	14.82	0.72	1.9	N	E520
64	08:51:18.20	+11:50:15.3	2.5	11(4)	2(1)	0.94 ± 0.28	−1.00 ± 0.19	41	104	14.41	0.61	4.8	Y	E1723
65	08:51:42.75	+11:34:08.8	2.5	10(4)	5(2)	0.81 ± 0.26	−1.00 ± 0.18					1.6	n	E3280; AGN2,z~0.3
66 ^a	08:52:10.06	+11:55:28.1	1.6	10(5)	8(4)	0.45 ± 0.55	0.03 ± 0.59	33		17.24	1.73	1.1	n	E721; Fore* (~M7V)
67	08:51:25.87	+11:59:33.2	1.7	9(4)	5(2)	1.00 ± 1.38	0.68 ± 0.31						n	AGN2?
68	08:51:25.06	+11:57:48.4	4.2	9(4)	3(2)	0.50 ± 0.44	−0.80 ± 0.48							
69 ^a	08:51:26.29	+11:37:06.3	2.6	9(3)	1(0)	1.00 ± 0.51	0.46 ± 0.40							
70 ^a	08:50:56.52	+11:45:21.9	2.8	7(4)	1(0)	0.54 ± 0.57	−0.34 ± 0.61	24	53	20.27	0.34	1.1	N	E362
71	08:52:06.51	+11:53:24.4	2.9	4(3)	1(1)	1.00 ± 0.29	−1.00 ± 0.45							
72 ^a	08:52:17.03	+11:41:11.3	2.9	4(3)	1(1)	1.00 ± 0.75	−0.37 ± 0.65	56		16.85	1.81	2.1	N	E207
New sources detected in Obs ID 0212080601														
73	08:50:36.96	+11:54:05.8	1.6	31(9)	19(6)	0.66 ± 0.20	−0.48 ± 0.33	35		14.49	0.78	1.2	Y	E1475
74 ^a	08:51:28.80	+11:48:07.9	1.7	22(7)	4(1)	1.00 ± 0.13	−0.74 ± 0.28		41					
75	08:50:58.56	+11:48:05.0	1.9	17(9)	4(2)	−0.26 ± 0.59	−0.14 ± 0.92			15.81	0.80	2.2	Y	E4630

et al. (2005).² 14 of our X-ray sources do not have an optical counterpart.

3.3 Members of M67 and source classification

There are several published proper motion membership studies for M67. Sanders (1977), Girard et al. (1989), and Zhao et al. (1993) used photographic plates and are limited to stars with $V \lesssim 16$. Yadav et al. (2008) calculated the proper motion membership probabilities for ~2400 stars towards M67 having $V \lesssim 22.5$. Recently, Vereshchagin et al. (2014) presented a revised list of M67 members among Yadav et al. sources using the convergent point method. We cross-matched our X-ray sources with Yadav et al. catalogue and found 38 matches. For these sources, we obtained the membership information from Vereshchagin et al., thus finding 19 members and 19 non-members. Among the sources lacking a counterpart in Yadav et al., we searched for membership information in Girard et al. to get two additional members. Note that for sources present in both, Yadav et al. and Girard et al., we give precedence to the former. For sources listed as non-members in Vereshchagin et al. (2014) that have ≥ 75 per cent probability in both, Yadav et al. and Girard et al. catalogues, we used the binarity and photometric information of the optical counterparts to argue membership, wherever applicable. Two X-ray sources, NX24 (Y892) and NX42 (EU Cnc), found to be non-members in Vereshchagin et al. (2014) and Yadav et al., have been reclassified as members of M67 based on archival photometric information and evidence from the literature. In total, we have detected 25 X-ray members of M67.

We classified the remaining 35 sources as AGN, quasar, galaxy, or foreground/background stars based on multiwavelength spectral energy distributions using photometry from the SDSS, WISE, 2MASS and NOMAD catalogues. The membership information for sources having an optical counterpart is given in column (13) of Table 2.

Since the membership of X-ray sources found by Belloni et al. (1998) and van den Berg et al. (2004) was based primarily on the work by Girard et al. (1989), we considered cross-matching their X-ray catalogues with Yadav et al. (2008), and looking at the membership of the corresponding source as given by Vereshchagin et al. (2014). Among the X-ray sources from Belloni et al. stated as non-members or sources without membership probability, we accept

four (RX47=NX17=CX7; RX17=NX22=CX15; RX42=CX17; RX35=NX73) as members, and among their proposed members, we reject four (RX43=NX11=CX20; RX44=NX63=CX68 is a galaxy; RX49=NX12=CX21 is a type-2 AGN; RX19). Note that CX20 (quasar) and CX68 were already classified as non-members by van den Berg et al. (2004). One source having unknown membership probability (RX54=NX62) in Belloni et al., has been classified as a member based on new proper motion information. Similarly, RX23 is found to be a non-member.

Among the *Chandra* X-ray sources (van den Berg et al. 2004) classified as probable non-members, we accept the following three as members based on their proper motion: (1) CX7=NX17=RX47 listed as a W UMa-type binary in SIMBAD; (2) CX15=NX22=RX17 is a known binary with circularized 1.2-day orbit; and (3) the unclassified source CX17=NX27=RX42). 12 sources close to the M67 main sequence were accepted as probable members by van den Berg et al.. Among these, we accept eight sources as true members of M67. These are: (1) CX58=NX30 is listed as RS CVn in SIMBAD, and a known binary with a 3.6-day period, (2) CX61=NX53 with 75 per cent proper motion membership probability in Yadav et al. (2008), and listed as a non-member in Vereshchagin et al. (2014), is a W UMa-type binary with period 0.27 d, has a photometric parallax distance of ~820 pc, and lies along the M67 main sequence, (3–5) CX62, CX77, and CX82 are members in Yadav et al. and Vereshchagin et al. with a high probability, (6–7) CX73, an M3V star, and CX76, a K7V star have estimated distances of ~950 pc based on SDSS and WISE photometry. CX73 has a counterpart in Yadav et al., where it has a proper motion membership probability of 48 per cent, but it is listed as a non-member in Vereshchagin et al. (2014). However, the hardness ratios of this source in the *Chandra* source catalogue (Evans et al. 2010) are consistent with that of an active star. CX76 has no counterpart in Yadav et al., (8) CX80 is a member of M67 in Yadav et al. and Vereshchagin et al.. Compared to the typical uncertainty in the *Chandra* X-ray source positions, the optical counterparts of CX73 and CX80 are quite far away (~4 arcsec). This could be attributed to the large off-axis distance of the location where these sources were detected on the ACIS CCD. We reject the membership of the remaining four sources (CX117, CX129, CX141, and CX153) due to their very small membership probabilities in Yadav et al., and their classification as non-members in Vereshchagin et al..

A compilation of the spectral and orbital parameters and X-ray luminosities for all the 43 members of M67 detected as X-ray sources

² <http://www.sdss.org/dr4/algorithms/sdssUBVRITransform.html>

Table 3. Summary of X-ray members of M67. The top panel shown members among the sources that were detected in *XMM-Newton* observations (this work) and classified as members in Vereshchagin et al. (2014), The middle panel shows members detected in *Chandra* van den Berg et al. (2004), while the lower panel lists members from *ROSAT* Belloni et al. (1998). The source IDs from this work, from Belloni et al. and van den Berg et al. are given in columns 1–3, the optical counterpart and its spectral type are listed in columns 4–5, X-ray luminosity in column 6, the orbital period and ellipticity in columns 7–8, and source classification is in column 9. See Section 3.3 for details.

(1) NX	(2) RX	(3) CX	(4) opt	(5) SpT	(6) $L_X/10^{30}$ (erg s $^{-1}$)	(7) P_{orb} (d)	(8) e	(9) Comments
Members from this work								
1	11	5	E1740	G0V	$6.7^{+1.1}_{-1.0}$	1.36	0	RS CVn
4	8	1	HU Cnc	G8IV	$5.6^{+1.3}_{-1.1}$	18.40	0.21	SGB, RS CVn
5	10	6	Y1289	G4III	$3.9^{+1.3}_{-2.1}$	42.83	0	III+WD,YS
6	13	9	HW Cnc	G4V	$6.7^{+1.5}_{-1.4}$	10.06	0	RS CVn
8	7	10	HT Cnc	F9V	$6.7^{+1.8}_{-1.7}$	1.36	0.10	RS CVn
10	40	16	AH Cnc	F7V	1.8 ± 0.14	0.36		Ecl W UMa
16	37	24	Y1476	G3III-IV	1.5 ± 0.14	1495	0.32	SB,YS
17	47	7	LN Cnc		2.7 ± 0.4	0.54		Ecl W UMa
20	45	19	EV Cnc	F3	1.10 ± 0.14	0.44		W UMa
21		23	HS Cnc	F9V	1.65 ± 0.27	0.36		W UMa
22	17	15	E1175	G9V	1.10 ± 0.14	1.17	0	RS CVn
27	42	17	E2759	G6V	0.96 ± 0.14			RS CVn?
29	38	48	HX Cnc	G1V	1.65 ± 0.41	2.66	0	RS CVn
30		58	HR Cnc		1.51 ± 0.27	3.58		RS CVn
33	50	49	E1784	G1V	0.55 ± 0.14	31.78	0.66	SB
36	53	36	E1581	F9V	1.37 ± 0.69	4.36	0.0	RS CVn, triple
42	16	57	EU Cnc	M5V	0.41 ± 0.14	0.09	0	CV, AM Her-type
45	46		Y1212	G0V	0.41 ± 0.14			
62	54		E1498		0.41 ± 0.14			
64	41	104	E1723	F9V	0.27 ± 0.14	var		RS CVn?
73	35		E1475	G8V	3.06 ± 0.97			CV?
75			E4630	~K5V	0.65 ± 0.32			
Members from Belloni et al. (1998)								
37	4	3	ES Cnc	F5IV	1.78 ± 0.69	1.07		Algol
	26		AG Cnc		6.1	2.82	0	RS CVn
	28		HY Cnc	G6V	2.1	2.65		RS CVn
	46	111	E1589	F7IV	0.08	7.16	0	RS CVn
24	52	47	Y892	G8III-IV	1.10 ± 0.14	698	0.11	RGB,YS
Members from van den Berg et al. (2004)								
53		61	ET Cnc	~G4V	0.69 ± 0.14	0.27		WU
		62	E683		0.40			
		67	E1781	G4V	0.36			binary
		72	E1677	G0V	0.16	5.7		RS CVn
		73	E2650	~M3V	0.46			
		76	E1720	~K7V	0.21			binary
		77	E394		0.23			binary
		78	E1730	F9V	0.30	5.95	0	RS CVn
		80	E2650	~M3V	0.20			
		81	E429	G5V	0.12	6.7		RS CVn
		82	E1208		0.16			
		88	E1729		0.09	7.65	0	RS CVn
		94	E1777	F8V	0.16			binary
		95	E1590	F6V	0.12	4913	0.34	BS
		155	E2983	F7V	0.04	11.02	0.26	SB,RS CVn
		157	HV Cnc	~F0V+M1V	0.02	10.34	0	Ecl Algol, triple

Notes: (a) The count rates and hence the luminosity values from Belloni et al. (1998) and van den Berg et al. (2004) do not have associated uncertainties. (b) Some of the known binaries but do not have an orbital/period solution. (c) The orbital parameters are from Mathieu (1990), Latham et al. (1992), van den Berg et al. (2000), and from unpublished work by D. Latham, R. Mathieu et al.

to date is given in Table 3. Here, we list the optical counterparts, the spectral types (from SIMBAD, VizieR, van den Berg et al. 2000, or estimated from archival multiwavelength photometry), the X-ray luminosities (from our work, van den Berg et al. 2004, or

Belloni et al. 1998, in that order of preference), the orbital periods and ellipticities (from van den Berg et al. 2000, 2004, or VizieR), and the source types (from van den Berg et al. 2004 or SIMBAD). Note that the luminosities for sources detected in the *XMM-Newton*

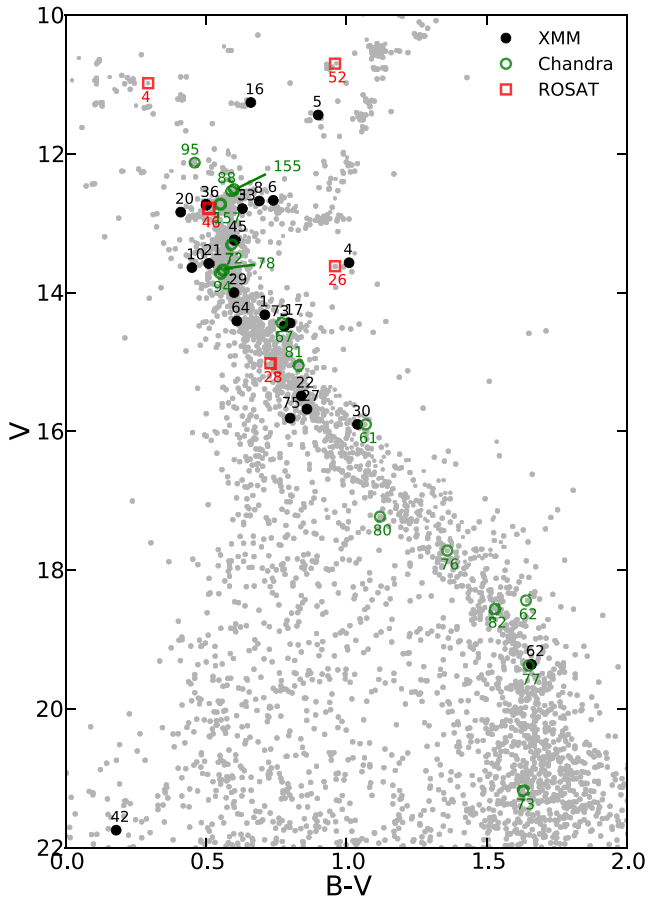


Figure 3. $V/(B - V)$ colour-magnitude diagram of X-ray members in M67 known till date, and listed in Table 3. Members from this work, from Belloni et al. (1998) and those from van den Berg et al. (2004) are plotted as black filled circles, red squares and green open circles, respectively. Numbers specified in these respective colours are the IDs of sources in the respective catalogues (NX, RX, and CX). Sources from WEBDA, from Yadav et al. (2008) with membership probability greater than 80 per cent, and all EIS sources are plotted in grey. Note that the $B - V$ colours for main sequence stars of spectral types A0, F0, G0, K0, M0, and M8 are 0.0, 0.3, 0.6, 0.8, 1.4, and 2.0, respectively (Johnson 1966).

observations are quoted for the 0.2–7 keV energy band, and those from van den Berg et al. are for the 0.3–7 keV energy band, while luminosities of Belloni et al. sources were calculated using the *ROSAT* PSPC count rate converted to flux in the 0.2–7 keV energy band assuming an APEC 1.5 keV plasma model. The orbital periods and eccentricities have been reproduced from van den Berg et al. (2000, 2004). For all known X-ray members of M67, we plot the $V/(B - V)$ colour-magnitude diagram in Fig. 3.

3.4 Chance identification

We estimated the probability of identifying an *XMM-Newton* source with an optical counterpart by chance using a procedure similar to the one adopted by van den Berg et al. (2004). This chance identification probability is a function of the positional uncertainty of the X-ray sources (Δ) and the surface density of optical sources in the M67 field. We calculated the mean positional uncertainty of the *XMM-Newton* sources, $\Delta = 1.6$ arcsec, and the number of EIS sources inside the fov, $N \approx 3035$. We used a circle with radius equal to the quadratic sum of Δ and the error in the optical positions,

1 arcsec, to calculate the search area around an X-ray source. Therefore, the probability that a randomly placed error circle includes an EIS-source is $N\pi(\Delta^2 + 1 \text{ arcsec}^2)/A = 1.1 \times 10^{-2}$. Here, the area of the fov (red curve in Fig. 1) is denoted as A approximated by a circle with radius 16 arcmin. Thus, for the 75 X-ray sources that we have considered here (Table 2), the probability of 0, 1 or 2 chance identifications with EIS-sources is 42, 37, and 16 per cent, respectively. Similarly, we calculated the probability for the chance identification of our X-ray sources with M67 proper motion members to be 87, 12, and 1 per cent, respectively. We therefore conclude that one or two *XMM-Newton* sources have been falsely identified with optical counterparts while our M67 X-ray members are all likely to be genuine. Another way of finding the chance identification probability is to note that among the 50 sources in van den Berg et al. (2004) having X-ray fluxes greater than the flux threshold of the *XMM* observations ($\sim 6 \times 10^{-15} \text{ erg cm}^{-2} \text{ s}^{-1}$, corresponding to a *Chandra* count rate of 1 count ks^{-1}), about 40 have an optical counterpart. We can then ask the question: given the probability of finding an optical source counterpart for an X-ray source as 0.8, what is the probability of finding counterparts to 62 *XMM* sources out of 75? Considering binomial distribution for the fraction of optical counterparts among X-ray sources, we get the answer as 10 per cent. The most likely value is 60 counterparts with a probability of 11.5 per cent.

We also estimated the probability for chance identification of *XMM-Newton* sources with *ROSAT* and *Chandra* sources. There are 47 *ROSAT* sources in the *XMM-Newton* fov, having an average error radius (90 per cent probability) as 12.25 arcsec, so the chance identification of 0, 1, and 2 counterparts is 56, 33, and 9 per cent, respectively. For the *Chandra* sources, we calculated the average error in RA and Dec separately, combined them with the corresponding average 1σ positional error for *XMM-Newton* sources, and then used twice the resultant errors in RA and Dec as the sides of the error box. Thus, considering the 154 *Chandra* sources in *XMM-Newton* fov we obtain the probability for 0, 1, and 2 chance identification as 97, 3, and 0 per cent, respectively. Accordingly, the probability of identifying *XMM-Newton* X-ray members with *ROSAT* and *Chandra* sources by chance is < 1 per cent.

3.5 Background sources

We used the soft AGN number counts from Gilli, Comastri & Hasinger (2007) to estimate the expected background sources in our observations. Using the flux limit of $5 \times 10^{-15} \text{ erg cm}^{-2} \text{ s}^{-1}$, we find from fig. 9 of Gilli et al. (2007) that there are ~ 230 AGN per deg^2 in the 0.5–2 keV band. Then, among the 75 unique sources in the *XMM* fov (16 arcmin radius circle), we expect ~ 50 AGN. Among the 75 unique sources in our X-ray source list given in Table 2, 24 are probable M67 members (Table 3), two are foreground stars. Four other non-members have V -band magnitudes less than 17, and are probably foreground stars. The remaining 45 sources have $V > 20$ or have no optical counterpart down to ~ 22 nd magnitude. These are all very likely to be AGN. In fact, we have confirmed that 18 of these are either galaxies (AGN status unknown) or AGN (see column 14 of Table 2).

4 SPECTRAL AND TIMING ANALYSIS

4.1 Spectral fitting and luminosities

We extracted PN background and source spectra for eight sources (NX1–8) (Table 2) having total counts > 100 . The spectrum for one of the sources (NX6) had to be extracted from the MOS1

Table 4. The best-fitting parameters and goodness-of-fit for the spectral fitting performed on sources NX1–8. The hydrogen column was held fixed and was calculated as described in Section 4.1. The columns are: (1) source ID; (2) single-multi-component model used for fitting along with photoelectric absorption; (3) and (4) temperatures of the two APEC components used to fit spectra of member stars of M67, (5) ratio of the emission measures of the two APEC components; (6) and (7) power-law index and blackbody temperature used for fitting non-members; (8) hydrogen column supplied to the photoelectric absorption model; (9) resulting fluxes along with their 90 per cent uncertainties; (10) energy conversion factor (ECF=Rate/Flux; see Section 4.1); and (11) the reduced χ^2 of the fit and the degrees of freedom (dof; equal to the number of channel groups minus one). The corresponding fits are shown in Fig. 4. Note that NX6 fitting was done on MOS1.

(1) NX	(2) Model	(3) kT_1 (keV)	(4) kT_2 (keV)	(5) EM1/EM2	(6) Γ	(7) kT (keV)	(8) N_H (10^{20} cm^{-2})	(9) $F/10^{-14}$ ($\text{erg cm}^{-2} \text{ s}^{-1}$)	(10) ECF/ 10^{11} ($\text{cts cm}^2 \text{ erg}^{-1}$)	(11) $\chi^2_{\nu}(\text{dof})$
1	2T	$0.39^{+0.66}_{-0.10}$	$1.95^{+0.90}_{-0.48}$	0.17 ± 0.10			2.2	$7.73^{+1.22}_{-1.18}$	6.6	0.82(14)
2	P+B				$1.36^{+0.62}_{-0.77}$	$0.19^{+0.04}_{-0.04}$	3.3	$10.83^{+1.70}_{-2.85}$	4.9	0.81(14)
3	P				$1.66^{+0.26}_{-0.25}$		3.2	$19.95^{+4.89}_{-3.67}$	4.7	0.69(8)
4	2T	$0.32^{+0.20}_{-0.08}$	$2.05^{+1.50}_{-0.51}$	0.24 ± 0.14			2.2	$6.45^{+1.50}_{-1.22}$	5.7	0.95(8)
5	2T	$0.76^{+0.16}_{-0.18}$	$3.30^{+0}_{-1.86}$	0.46 ± 0.29			2.2	$4.55^{+1.54}_{-2.42}$	6.2	0.63(5)
6	2T	$0.24^{+0.24}_{-0.12}$	$1.38^{+0.42}_{-0.15}$	0.43 ± 0.30			2.2	$7.71^{+1.73}_{-1.61}$	1.8	1.08(3)
7	P				$1.71^{+0.38}_{-0.35}$		3.3	$6.39^{+2.02}_{-1.41}$	4.1	0.91(7)
8	T		$1.51^{+0.55}_{-0.34}$				2.2	$7.71^{+2.06}_{-2.00}$	6.9	0.94(7)

detector since it is located in the PN CCD gap. To find the appropriate regions for the extraction of background and source spectra, we used DS9. Spectral information was extracted from within a circle having 30 arcsec radius centred on the X-ray source. Background spectra were extracted from annuli with inner radii between 35 arcsec and 45 arcsec and widths between 60 arcsec and 75 arcsec centred on the X-ray sources, taking care to avoid contamination from nearby sources. For those X-ray sources where the annulus would spill over to the neighbouring CCD, the region of choice for background spectral extraction was a circle close to and having a similar RAWY as for the source extraction region, and with radius between 40 arcsec and 50 arcsec. We used the *SAS* task *especget* to perform the extraction of source and background spectra for NX1–8. We grouped the spectral (PHA) channels using the *specgroup* task such that each grouped PHA channel had 16 counts, sufficient for the assumption of Gaussian distribution of uncertainties. This task also performs the subtraction of background spectrum from the source spectrum.

The spectral fitting was done using the *XSPEC* package. We ignored spectral groups containing channels with energies < 0.2 keV and > 12 keV. X-ray spectra for the M67 members were fitted using an absorber model (*phabs*) multiplied to a model of collisionally-ionized plasma (*apec* with one or two characteristic plasma temperatures and both having default elemental abundances). The column density N_H for the absorber was fixed at $2.2 \times 10^{20} \text{ cm}^{-2}$ (converted from $E(B - V) = 0.04$ using $R_V = 3.1$). Note that the choice of two *apec* components instead of a single one was decided based on the reduced χ^2 of the fit and manual inspection of the fitted spectrum and residuals. For all the other X-ray sources, a photoelectrically absorbed power law with or without a blackbody component, as expected for AGN, was used for spectral fitting. In this case, the N_H was estimated using the *nh* task. For NX1–8, the fitted model, their best-fitting parameters with 90 per cent uncertainties given by *XSPEC*, and the N_H used are shown in columns 2–8 in Table 4, and the resulting reduced χ^2 and degrees of freedom (dof) are in column 11. Fig. 4 shows the spectra for the eight sources considered for fitting, of which five are members of M67.

The fluxes of NX1–8 in the 0.2–7 keV energy band were calculated using the best-fitting parameters in the *flux* task in *XSPEC*. The resulting fluxes along with their 90 per cent uncertainties are

given in column 9 of Table 4. The energy conversion factors (ECF=Rate/Flux) are tabulated in column 10. For members of M67 among sources in Table 4, we find that the mean ECF is $6.3 \pm 0.4 \times 10^{11} \text{ counts cm}^2 \text{ erg}^{-1}$ (uncertainty in the count rate not folded in). We used this factor to calculate the X-ray fluxes in the 0.2–7 keV band for all the M67 members detected in the *XMM* observations. Similarly, we determined the ECFs for the 0.3–7 keV and 0.1–2.4 keV bands. The former band is used by van den Berg et al. (2004) for *Chandra* observations of M67 while the latter is relevant for *ROSAT*. For coronal sources, we used the ECFs to calculate the multiplicative factors for converting from the *Chandra* and *ROSAT* bands to the 0.2–7 keV band as 1.0 ± 0.1 and 0.8 ± 0.1 , respectively. X-ray luminosities in the 0.2–7.0 keV energy band of *XMM* for all the members of M67 calculated assuming a mean distance of 850 pc are given in Table 3. For Obs ID 0212080601, we fitted the spectra for four bright members of M67 (NX4–6, NX8) using the parameters listed in Table 4 as initial values, and calculated the ECFs. The mean ECF was found to be $6.2 \times 10^{11} \text{ counts cm}^2 \text{ erg}^{-1}$. We used this ECF to convert the count rates of NX73–75 to their corresponding fluxes.

4.2 Hardness ratios

Spectral information for weak X-ray sources can be obtained by comparing the count rates in coarse spectral bins after applying the correction for exposure, vignetting, and background. Taking advantage of the wide bandpass of *XMM*, we defined two hardness ratios, HR1 and HR2, as given in equation (1).

$$\text{HR1} = \frac{(M_X - S_X)}{(M_X + S_X)}, \quad \text{HR2} = \frac{(H_X - M_X)}{(H_X + M_X)}, \quad (1)$$

where S_X , M_X , and H_X denote the exposure-corrected counts in the soft band (0.2–0.5 keV), the medium band (0.5–2.0 keV), and the hard band (2.0–7.0 keV), respectively. For each X-ray source detected, HR1 and HR2 are calculated within the best-fitting source position, using the best-fitting count rates in all the bands by the source-finding algorithm *emldetect*. They are tabulated in column 7 of Table 2 for the PN detector (unless noted as MOS1). Sources with very soft spectra will have $\text{HR1} < 0$ and those with hard spectra will have $\text{HR2} > 0$. In Fig. 5, we plot HR2 against HR1 for all of our

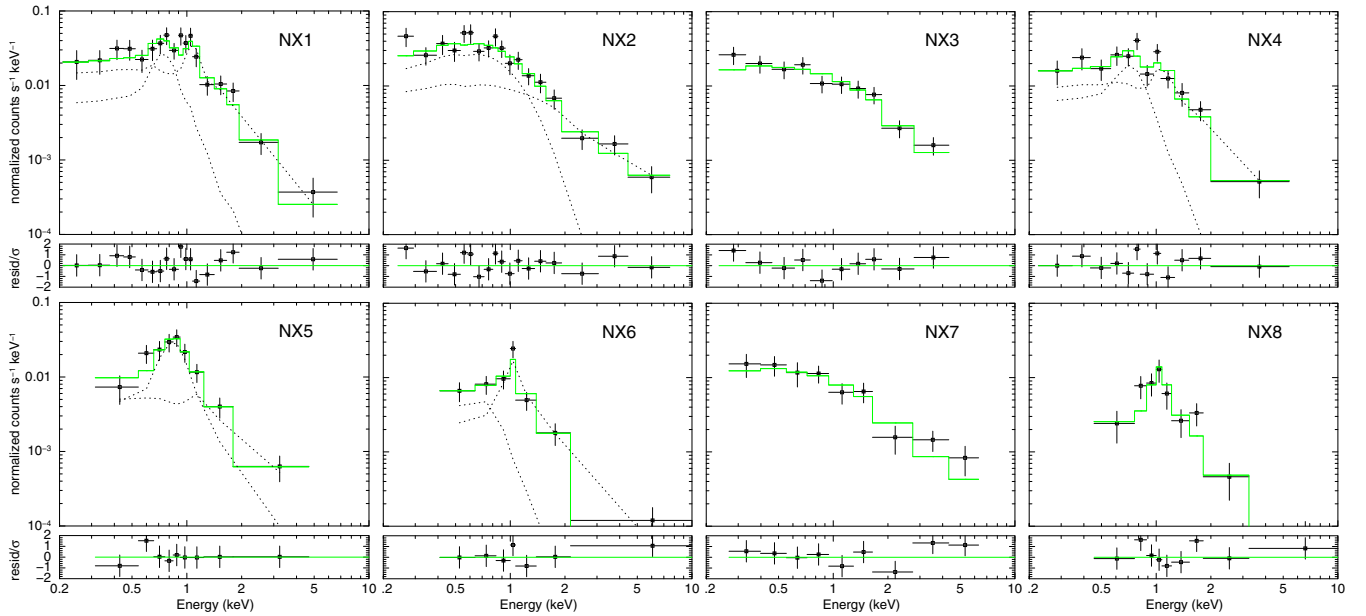


Figure 4. Spectral fitting for sources having PN count rate greater than 100 (NX1–8). For each source, the top panel shows the binned background-subtracted normalized count rate (black points with error bars), the best-fitting single-/multicomponent spectrum (green), and the individual spectral components (dashed black curves), while the bottom panel shows the residual-to-noise ratio. For best-fitting parameters, see Table 4.

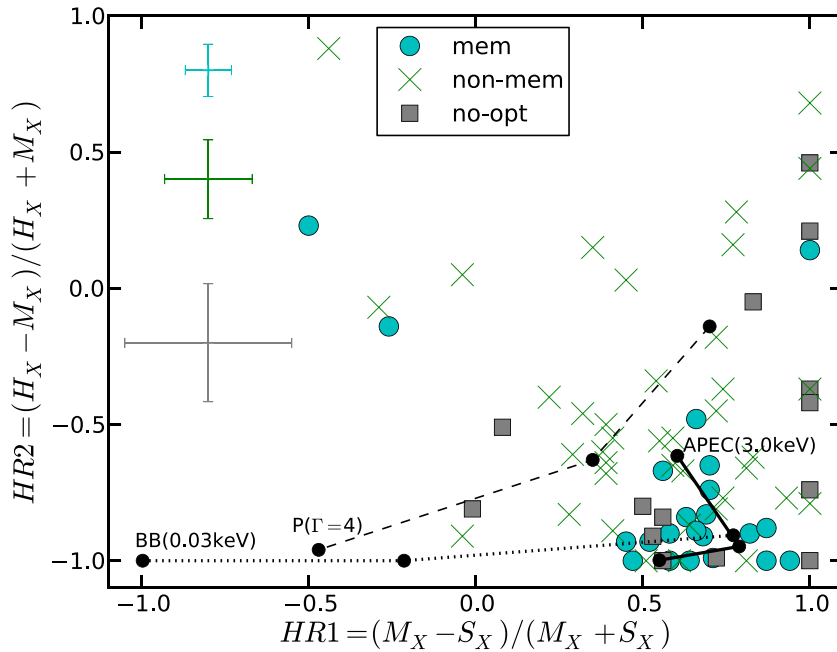


Figure 5. Hardness ratios (see Section 4.2) of the 75 unique sources detected in the *XMM* observations. Sources identified as M67 members, non-members, and sources with no optical counterparts (also suggested as non-members in Section 3.3) are shown as cyan circles, green crosses, and grey squares, respectively. The weighted mean error bars are shown to the left-hand side of the plot, representative of sources in each of the three classes. Note that these error bars are just for guiding the eye, and the actual uncertainties (see Table 2) depend on count rate. Note also that the members NX36 and NX75 have large uncertainties in their hardness ratios (~ 0.6), and appear to be outliers with respect to the clustering of M67 members in the lower right-hand corner of the plot denoting coronal emitters. The cataclysmic variable EU Cnc is located at $HR1 = -0.50 (\pm 0.33)$, $HR2 = 0.23 (\pm 0.46)$. The expected hardness ratios for three kinds of spectra, power-law ($P(\Gamma)$), blackbody (BB), and collisionally-ionized plasma (APEC), are shown as black circles joined by black lines (dashed, dotted, and solid lines, respectively). Each black dot marks a unique value of temperature (in case of BB or APEC) or a power-law index. The hardness ratios have been calculated for $\Gamma=1, 2$, and 4 in the case of power law; $T=0.03, 0.1$, and 0.3 keV for BB; and $T=0.3, 1$, and 3 keV for APEC. For each type of spectrum, one point has been labelled. These were calculated using WebPIMMS using Galactic hydrogen column density of $2.2 \times 10^{20} \text{ cm}^{-2}$.

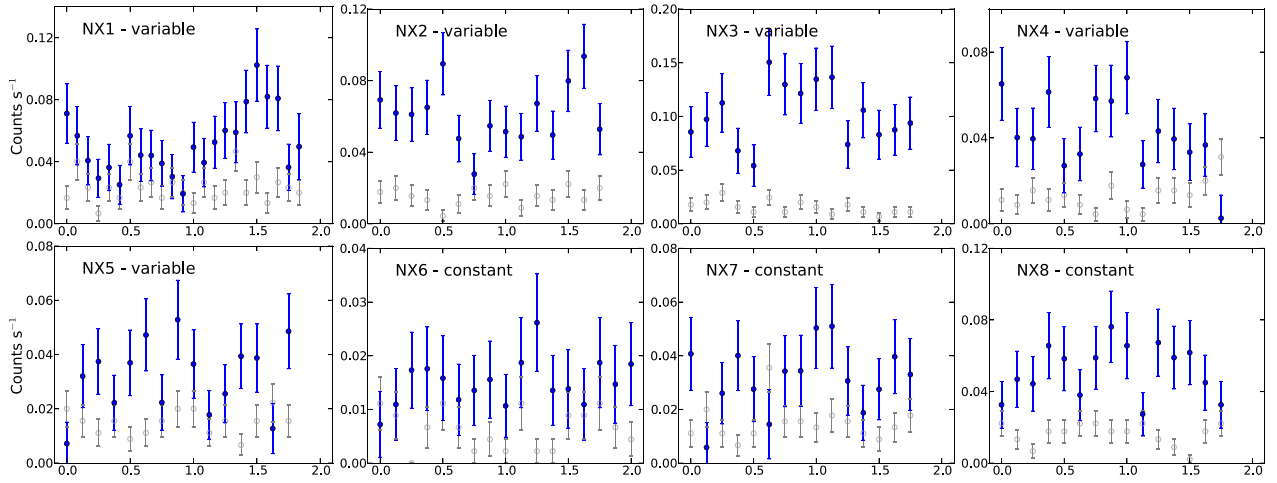


Figure 6. The light curves for sources having 100 or more counts in PN (MOS1 for NX6). The background-subtracted light curves are in blue and the background-only light curves are in grey. Count rate (counts per second) is on y-axis and time in hours is along the x-axis. The time bin size is set to 300 s for NX1 and 450 s for the rest in order to get >30 counts in each bin. Whether a source is variable or not according to the χ^2 test is denoted on the upper-left side. See Section 4.3 for details.

75 unique sources. Apart from NX36 and NX75, whose hardness ratios are imprecise, and the cataclysmic variable NX42=EU Cnc, all members of M67 lie in the region $HR1 \gtrsim 0.5$, $HR2 \lesssim 0.5$. AGN and other non-members also tend to lie in this region, but their scatter is much larger. Hardness ratios expected from power-law (P), blackbody (BB) and collisionally ionized plasma (APEC) are also shown. These were calculated using *WEBPIMMS* using Galactic hydrogen column density of $2.2 \times 10^{20} \text{ cm}^{-2}$. Note that the uncertainties of the source hardness ratios depend on the count rate and for our *XMM* sources they have a large scatter. Only the weighted mean uncertainties are shown as guidance in Fig. 5 for each class of object.

4.3 X-ray variability

The X-ray light curves were extracted from the same regions as for the spectra from the filtered PN (MOS1 for NX6) event list in the 0.2–10 keV band with tasks *evselect* and *epic1ccorr*. The time bin size was set to 300 s (NX1) or 450 s (NX2–8) to get adequate signal-to-noise ratio in each bin. For PN, we selected only those events with *PATTERN* keyword less than or equal to 4. The light curves for NX1–8 are shown in Fig. 6. The background-subtracted light curves are in blue and the background-only light curves are in grey.

We used the χ^2 test to check for the departure from white noise in the light curves, and selected sources having χ^2 probability <25 per cent as significant variables. Sources NX1 to NX5 are thus found to be variable, while the variability in NX6–8 is not significant. Members of M67 showing a significant variability are NX1, NX4, and NX5, which also exhibit short-term autocorrelation. For the RS CVn binary, NX1, the autocorrelation function shows a significant correlated structure for lag time $k \leq 2$ implying variability time-scale of ~ 10 min.

We also compared the X-ray luminosities of M67 members detected in the *XMM* observation ID 0109461001 with those in the *Chandra* observation from van den Berg et al. (2004) separated by five months, and between ID 0109461001 and ID 0212080601 to find the variability over two observations spaced by ~ 4 yr. The comparison plots are shown in Fig. 7.

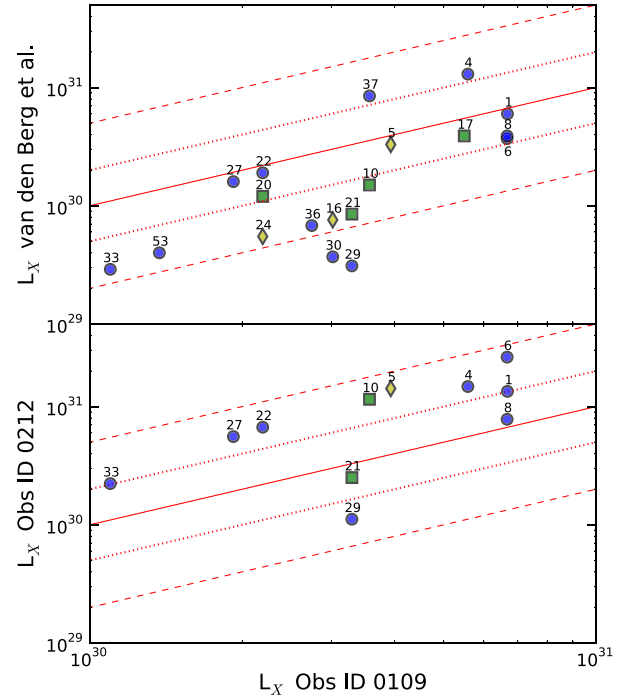


Figure 7. Variability of X-ray members of M67 on 5-month (top) and 4-yr (bottom) time-scale shown using plots of the X-ray luminosities of sources in Obs ID 0109461001 versus X-ray luminosities in van den Berg et al. (2004) (top) and those in Obs ID 0212080601 (bottom), respectively. The solid red line marks equal luminosities in either quantity plotted, and the red dotted (dashed) lines denote variability of a factor of 2 (5). RS CVn or Algol systems are marked with blue circles, W UMa systems with green squares, and all other M67 members (all of which are yellow stragglers) are shown as yellow diamonds. See Section 4.3 for details. Note that the implied X-ray luminosities in the 0.2–7.0 keV energy band are used.

The RS CVn-type systems, NX1, NX6, NX22, and NX27, have relatively stable X-ray emission between ID 0109461001 and van den Berg et al., but are variable by more than a factor of 2 between ID 0109461001 and ID 0212080601, while NX8 appears stable between all three observations. NX4, NX16, NX33, NX36, NX37, and

NX53 are variable by more than a factor of 2 between any of the two observations. NX29 and NX30 are variable by more than a factor of 5 over the observations spaced by five months. The W UMa-type system, NX21, varied by a factor of 4 between ID 0109461001 and *Chandra* observations. NX10 is variable by more than a factor of 2 over the observations separated by months as well as years. Both types of systems, RS CVn-type and W UMa-type, are known to be highly variable due to recurrent flaring activity. The snapshot observations, which are not very long, could have caught them in different phases of their activity. Three yellow stragglers, which are present in the multi-epoch data, also display average change in luminosity by a factor of >2 . Fig. 7 suggests that, the variability amplitudes of RS CVn-type sources, W UMa-type systems, and yellow stragglers are similar.

5 NOTES ON INDIVIDUAL CLASSES OF STELLAR X-RAY SOURCES AND THEIR X-RAY LUMINOSITY FUNCTIONS

In Table 5, we provide notes on the individual classes of stellar X-ray sources in M67. We have used the orbital parameters for X-ray members from van den Berg et al. (2004), who have in turn compiled them from Mathieu (1990), Latham et al. (1992), van den Berg et al. (2000), and from unpublished work by D. Latham, R. Mathieu et al. For the names of the sources and their parameters, please see Table 3.

In Fig. 8, we plot the X-ray luminosity functions (XLFs) of various stellar types in M67 and compare them with the XLFs of other open star clusters of intermediate-to-old age. The star clusters used for comparison with M67 (4 Gyr; Table 3) are: NGC 6633, IC 4756 (0.7 Gyr; Briggs et al. 2000), NGC 6791 (8 Gyr; van den Berg et al. 2013), and NGC 188 (7 Gyr; Belloni et al. 1998; Gondoin 2005). The clusters in general have different median values for the X-ray luminosity for different stellar types. The XLF of the RSCVns in M67 is shown in the top panel of Fig. 8 along with that in the other clusters. The XLF of field RS CVn binaries from Singh, Drake & White (1996) is also shown for guidance. In general, RS CVn-type systems are more numerous at the low-luminosity end, and there exist very few systems above $L_X = 10^{31}$ erg s $^{-1}$. It is evident that the highest luminosity of RS CVn-type systems seen in the clusters is a function of the age of the clusters. Older clusters appear to have larger number of brighter luminosity systems. In other words, the XLFs indicate that the RS CVn-type systems get brighter as they age. It is also evident from Fig. 8 that the RS CVn XLFs of M67 and the Singh et al. sample have a much shallower slope than those of the other clusters, which may be due to two distinct distributions of RS CVn-type systems in open clusters.³ This is suggestive of the intuitive notion that, as the clusters age, more active binary systems and active stars are produced (see also Verbunt 1999; van den Berg 2013). This picture is consistent with younger systems having lower luminosity and older ones being more luminous. The shallower slope of and the presence of several $L_X > 10^{31}$ erg s $^{-1}$ systems among the field population advocate that these are old systems which may have been kicked out of their parent clusters.

The second panel from the top of Fig. 8 shows the XLFs of WU UMa-type and FK Com-type systems in M67, NGC 188, and

NGC 6791. Down to the sensitivity of X-ray surveys in NGC 6633 and IC 4756, these clusters do not have any definite detections of contact or coalesced binaries. From Fig. 8, we see that the slopes of the contact binary XLFs are steep, and there are several systems having X-ray luminosities greater than 10^{30} erg s $^{-1}$. The WU UMa-type and FK Com-type systems in NGC 188 are more luminous ($L_X > 4 \times 10^{30}$ erg s $^{-1}$) compared to those in M67 and NGC 6791. This is surprising given that NGC 188 is neither the most massive (in fact, it is the least massive among the old clusters) nor the oldest cluster among the sample.

The third panel from the top of Fig. 8 shows the XLFs of stragglers and peculiar systems. The brightest ends of the XLFs of M67, NGC 188, and NGC 6791 are steep. Majority of the straggler and peculiar systems known in these clusters span a narrow range in X-ray luminosity, 3×10^{30} to 6×10^{30} erg s $^{-1}$. This could indicate a fine-tuning of X-ray emission for the brightest systems, which needs further investigation. At the low-luminosity end, incompleteness due to survey sensitivities is apparent. As such, among the old clusters, the XLFs appear to be independent of the cluster age.

The bottom panel of Fig. 8 shows the XLFs of cataclysmic variables. The XLF of a mixed sample of polars from a variety of binaries from Barrett, Singh & Mitchell (1999) is also shown for guidance. NGC 6633, IC 4756, which are 0.7 Gyr old, do not have any CVs. M67 (4 Gyr old) has only one CV, the polar in quiescence EU Cnc, which has 0.2–7 keV X-ray luminosity of $<10^{30}$ erg s $^{-1}$, while the older cluster NGC 6791 contains three CVs with $L_X \gg 10^{30}$ erg s $^{-1}$. NGC 188, although an old cluster of age 7 Gyr, does not have any CVs brighter than 10^{30} erg s $^{-1}$. Given this information, the production of bright CVs appears to be correlated with the mass of the cluster. NGC 6791 is about 10 times more massive than M67, and M67 is about twice as massive as NGC 188 (Pandey, Bhatt & Mahra 1987; van den Berg et al. 2013, and references therein). van den Berg et al. (2013) find that the number of CVs per unit mass is consistent with the field, and conclude that the three CVs which are confirmed members of NGC 6791 are primordial. Unlike RS CVn systems, in the case of CVs we cannot ascertain the evolution of the XLF with age. However, the mass of the parent cluster has to play an important role. More CVs are produced as clusters age, and it is plausible that these systems also get brighter with age. In this picture, the Barrett et al. (1999) sample, which contains numerous polars at high luminosity ($L_X > 10^{31}$ erg s $^{-1}$), could be very old systems. The shallow slope of the Barrett et al. XLF is likely due to the bimodal distribution of polars in X-ray luminosity as noted by Barrett et al., which may be related to age of the systems.

6 SUMMARY

Our analysis of two *XMM-Newton* observations of the old open cluster, M67 has led to the detection of 25 members of M67, of which 24 have been detected in previous observations with *Chandra* (van den Berg et al. 2004) and *ROSAT* (Belloni et al. 1993, 1998). One X-ray source is a newly-detected member. Based on present and older observations we have compiled an updated list of 43 X-ray members of M67 (Table 3), 16 of which are likely to be RS CVn or related binaries with circularized or near-circularized orbits having periods $P_{\text{orb}} \lesssim 10$ d. We also detected five contact binaries with $P_{\text{orb}} \lesssim 0.5$ d. X-ray emission from three yellow stragglers and two blue stragglers detected here is not well understood. Another peculiar source, HU Cnc, lies below the intersection of the subgiant and red giant branches in the HR diagram (Fig. 3), and may be a coronal emitter. The only cataclysmic variable known in M67,

³ We could have also missed a population of low-luminosity RS CVns due to the limit of our sensitivity, but this would only partly explain the shallow slope.

Table 5. Notes on individual classes of stellar X-Ray sources.

Class/Type	Source(s)	Notes
RS CVn ^a	(1) NX6, NX8, NX36	Lie along subgiant branch. $P_{\text{orb}} \lesssim 10$ d, $e \lesssim 0.1$.
	(2) NX1, NX22, NC29, RX26, RX46, CX78, CX88	Proposed as RS CVn systems by van den Berg et al. (2004) and Belloni et al. (1998) based on the presence of Ca H&K emission lines, X-ray spectrum, orbital period and circularisation. Lie along subgiant branch. $P_{\text{orb}} < 8$ d, $e \simeq 0$.
	(3) NX27, NX64, RX28	Candidate RS CVn-type systems or their main-sequence analogues. Complete orbital solution not available.
Algol	(1) NX37	Previously known Algol system. $P_{\text{orb}} = 1.1$ d.
	(2) CX157	Candidate Algol-type system. Consists of an eclipsing system containing an F0-type and an early M-type main sequence star (Gökay, Grol & Derman 2013).
Contact binaries	NX10, NX17, NX20, NX21, NX53=CX61	All are W UMa-type systems. $P_{\text{orb}} \lesssim 0.5$ d. NX10 and NX17 are eclipsing binaries. NX61 had been assigned unknown membership probability by van den Berg et al. (2004), and we confirm its M67 membership in this work. NX10, NX17, NX21, and NX53 are on the main sequence. NX20 is close to the main-sequence turnoff.
CV	(1) RX35=NX73	Belloni et al. (1998) use $B - V = 0.38$ from Sanders (1977) and hardness of the X-ray spectrum to claim that this could be an accreting white dwarf system. Orbital parameters are unknown. However, EIS has $B - V = 0.79$, and our hardness ratios are $\text{HR1} = 0.66 \pm 0.20$, $\text{HR2} = -0.48 \pm 0.30$. The position in the HR diagram and absence of soft and hard X-ray components therefore argue against the CV hypothesis for this source.
	(2) NX42	AM Her-type CV, EU Cnc. <i>ROSAT</i> detection is only below 0.4 keV (Belloni et al. 1998), but van den Berg et al. (2004) find a relatively high hardness ratio. In this work, we have $\text{HR1} = -0.50 \pm 0.33$ and $\text{HR2} = 0.23 \pm 0.46$, consistent with the high hardness ratios typical of magnetic CVs.
Blue Stragglers	(1) NX37	1.1 d orbit for the inner binary of a possible triple system (Goranskij et al. 1992; van den Berg et al. 2001). van den Berg et al. (2000) find the spectral type to be F5IV, and van den Berg et al. (2001) suggest that the X-ray emission could be due to magnetic activity in the rapidly rotating subgiant.
	(2) CX94 ^b	Noted as a blue straggler in Sandquist & Shetrone (2003), but with the EIS $B - V$ colour, it is very close to main-sequence turnoff. van den Berg et al. (2004) noted that CX94 shows no radial-velocity variations over six observations spanning 3923 d.
	(3) CX95 ^b	$P_{\text{orb}} = 4913$ d and $e = 0.3$. The primary has spectral type F6V and a projected rotation speed, $v \sin(i) \simeq 20 \text{ km s}^{-1}$ (Latham & Milone 1996; Liu, Deng & Chavez 2008). van den Berg et al. postulate that CX95 could have an undetected close binary, hence explaining the X-ray emission.
Yellow Stragglers	(1) NX5	Binary system with a G4 giant and a cool white dwarf, with a circular orbit having $P_{\text{orb}} = 43$ d. Ca II H,K emission lines are present Pasquini & Belloni (1998). Circularized orbit suggests strong tidal forces between the two components of the system, but the cool temperature of the white dwarf precludes mass transfer.
	(2) NX16, NX24	$P_{\text{orb}} = 1495$ d and 698 d, and $e = 0.32$ and 0.11, respectively. Tidal interaction therefore not strong enough to explain their X-ray emission, $L_X > 10^{30}$.
Peculiar System	NX4	Lies about one mag below the subgiant and giant branches in the $V/B - V$ diagram (Fig. 3). The X-ray spectrum and variability suggest that the X-ray emission is coronal.
New M67 Member	NX75	Associated with the $V=15$ mag optical counterpart, E4630. Orbital parameters and hardness ratios are not well constrained. We estimate the spectral type of this source as K4V and a distance of 860 pc based on archival multiwavelength photometry, and thus suggest that it may be an RS CVn type. Further optical spectroscopy and radial velocity measurements are needed for understanding its chromospheric activity and for deriving its orbital parameters.
Other Members with unknown class	(1) NX45, CX67, CX94	Lie along the F/G/K main sequence and are known binaries with undetermined orbital solutions. NX45 has a spectral type G0V and is close to the main-sequence turnoff.
	(2) NX62, CX62, CX73, CX77, CX80, CX82	Lie along the M dwarf part of the main sequence of M67. These members are possibly spun-up binaries with active chromospheres where both members are main sequence stars.

(a) Sources with $P_{\text{orb}} \lesssim 10$ d are expected to have near-circularized orbits (e.g. Latham 2007). (b) CX94 and CX95 are within the fov of the *XMM-Newton* observations, but below the detection threshold.

the AM Her-type system EU Cnc, is detected in the *XMM-Newton* observations considered in this work. 14 X-ray members do not have any definitive classification; some of these are known binaries with undetermined orbital solutions. Spectroscopic follow-up observations and radial velocity monitoring in the optical and at other wavelengths is required for elucidating the nature of these sources.

M67 lies within the field of *Kepler*'s K2 observing campaign,⁴ and the optical light curves from this campaign will be useful for calculating orbital parameters of at least some of these systems.

⁴ <http://keplerscience.arc.nasa.gov/K2/>

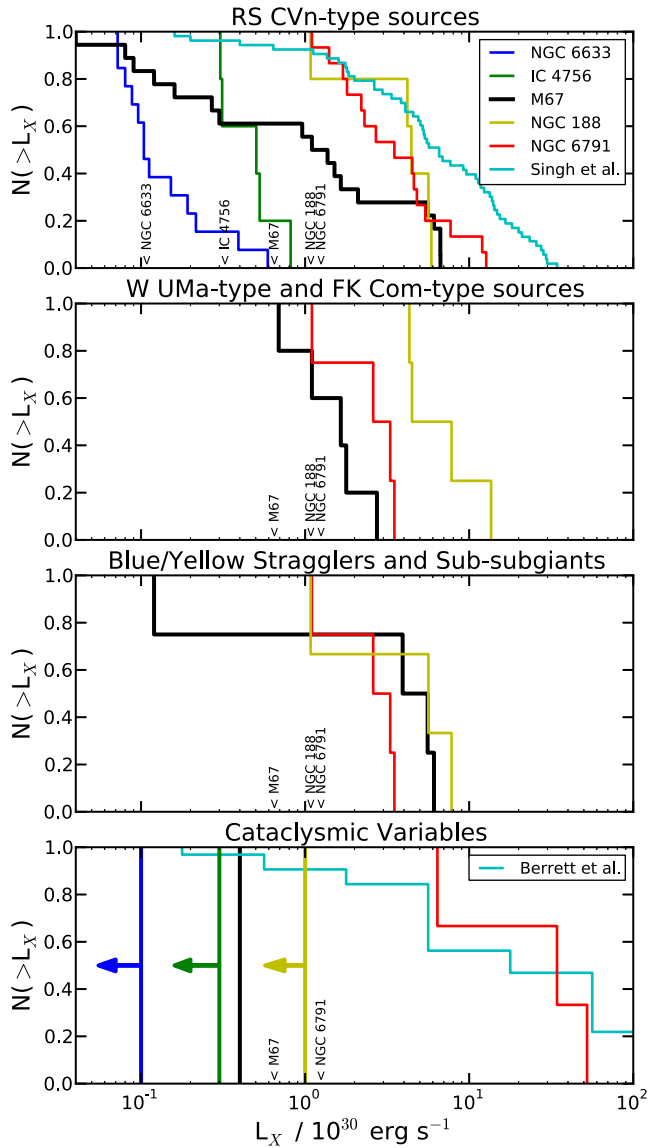


Figure 8. The XLFs for RS CVn-type sources (top panel), contact and coalesced binaries (second panel from top), stragglers and other sources having anomalous locations in the colour–magnitude diagram (third panel from top), and CVs (bottom panel) in different star clusters: NGC 6633, IC 4756 (0.7 Gyr), M67 (4 Gyr), NGC 188 (7 Gyr), and NGC 6791 (8 Gyr). The sensitivity threshold for the different star clusters below which the source counts are incomplete are shown towards the bottom of each panel. See Section 5 for details.

We already have an existing programme to study the ultraviolet emission from M67 members and to compare it with their X-ray emission (Sindhu, Subramaniam & Radha 2015). We found X-ray variability among M67 members on a five-month time-scale and an ~ 4 -yr time-scale by comparing the *XMM-Newton* observations and *Chandra* observations from van den Berg et al. (2004). The fractional variability of RS CVn-type sources, W UMa-type systems, and yellow stragglers are found to be quite similar, factor of ~ 3 on months and years time-scales. Finally, we have studied the XLF of RS CVn-type (and other types) members in M67. A comparison with that of other intermediate-to-old open clusters (Fig. 8) shows an increased number of higher luminosity stars in the older clusters, thus suggesting that more active binary systems and active stars are likely produced with the aging of clusters. Deeper and more

sensitive X-ray and optical observations of intermediate-to-old star clusters are encouraged in order to find fainter sources and thus extend the luminosity functions of the different classes of X-ray sources to the faint end.

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